Uncertainty in measurements of the distances of scattering screens in pulsar observations

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Abstract In our previous paper we investigated properties of the ionized interstellar medium in the direction of three distant pulsars: B1641–45, B1749–28 and B1933+16. We found that uniformly distributed scattering material cannot explain measured temporal and angular broadening. We applied a model for a thin scattering screen and found the distances to the scattering screens in all directions. In this paper, we consider more complicated models of scattering material distribution, such as models containing both a uniformly distributed medium and thin screen. Based on these models, we estimate the accuracy of localization of scattering screens and the possible relative contribution of each scattering component.

Key words: ISM: structure — pulsars: general — scattering

1 INTRODUCTION

Inhomogeneities in the interstellar plasma distort radio emissions, which leads to angular broadening of radio sources. The angular diameters of scatter-broadened high-latitude ($|b| > 10°$) extragalactic sources decrease with galactic latitude (Readhead & Hewish 1972), whereas for low-latitude sources ($|b| < 10°$) there is a tendency for them to be larger the closer they are to the Galaxy's center (Rao & Ananthakrishnan 1984). In contrast, Lazio et al. (2008) found that there is no correlation between the galactic latitude and the scattering diameter in the direction of the anticenter of the Galaxy that can be explained by a small quantity of uniformly distributed scattering material in the external part of the Galaxy.

Gwinn et al. (1993) showed that temporal and angular broadening of the radiation from pulsars corresponds to scattering by a uniform medium, except for several young pulsars which are scattered by their supernova remnants. The same conclusion was made by Britton et al. (1998). Stinebring et al. (2001) found parabolic structures in the secondary spectra (a secondary spectrum is the power spectrum of the dynamic spectrum) of some

pulsars that arise as a result of scattering by material situated in compact areas (known as thin screens) in the line of sight.

The ground-space radio interferometer RadioAstron has a sufficient angular resolution to resolve scattered discs of some pulsars, which allows us to study the distribution of scattering matter in detail (Gwinn et al. 2016; Popov et al. 2016, 2017; Shishov et al. 2017; Andrianov et al. 2017). The summary of main RadioAstron results was introduced in the All-Russian Astronomical Conference VAK-2017 that was held on 2017 September 17 (Samus & Li 2018).

2 MODELS OF THE DISTRIBUTION OF SCATTERING MATTER

2.1 Simple Scattering Models

In our previous paper (Popov et al. 2016), we measured independently both the angular diameter of the average scattering disk θ_H and the temporal broadening time τ_{sc}. Dependences of these values on mean scattering angle per unit length $\psi(z)$ have been given by expressions

(Blandford & Narayan 1985)

$$\theta_{\mathrm{H}}^2 = \frac{4\ln 2}{D^2} \int_0^D z^2 \psi(z) dz \,, \qquad (1)$$

$$\tau_{\mathrm{sc}} = \frac{1}{2cD} \int_0^D z\,(D-z)\,\psi(z) dz \,, \qquad (2)$$

where z is a coordinate along the line of sight from source ($z = 0$) to observer ($z = D$), and c is the speed of light. Gwinn et al. (1993) (see also Britton et al. 1998) have considered two important cases: uniformly distributed scattering material ($\psi(z) = \Psi_0$) and scattering material concentrated in the thin screen ($\psi(z) = \Psi_1 \delta(z - (D - d_s))$). Here d_{s} is the distance from the observer to the screen, and Ψ_0 and Ψ_1 are constants. Following this approach, we have shown that measured θ_{H} and τ_{sc} do not satisfy the relationship for a uniform medium which is $\theta_u^2 = 16\ln 2 c\tau_{\mathrm{sc}}/D$. Using the thin screen model, we have calculated screen distances for three pulsars.

It is clear that the interstellar medium has a more complex composition than a thin single layer on the way to a pulsar. Structures that scatter radio waves might be localized mainly into areas with high turbulence, such as shells of H II regions, whose sizes usually are only a small part of the distance to the pulsars. Several ionized clouds with a different scattering power may be located along the line of sight. Uniformly distributed interstellar medium in the long distances also may produce significant scattering. The same temporal and angular broadening can be produced not only by a single screen but also by several thin screens with different scattering power or by a mixture of one or more screens, immersed in the uniform medium. In this regard, it seems important to consider more sophisticated models of the interstellar medium to figure out if there is a dominant scattering screen. If such a screen exists, we need to assess the accuracy of localization for this screen under the assumption that other scattering agents exist.

2.2 Thin Screen and Uniform Medium

Gwinn et al. (1993) proposed a two-component model that includes scattering by both uniformly distributed material and a thin screen. For such a model,

$$\psi(z) = \Psi_1 \delta(z - (D - d_{\mathrm{s}})) + \Psi_0 \,.$$

In this case, the relationship between θ_{H} and τ_{sc} is as follows

$$\theta_{\mathrm{H}}^2 = \theta_u^2 \frac{\chi + 3s^2}{\chi + 6s\,(1-s)}, \qquad (3)$$

where $\chi = \Psi_0 D/\Psi_1$ is the relative power of scattering and $s = (D - d_{\mathrm{s}})/D$ is the fractional distance of the screen from the pulsar to the observer. When all scatterings are produced by the screen ($\Psi_0 D \ll \Psi_1$ and $\chi \to 0$), we obtain the expression describing the thin screen alone, $\theta_{\mathrm{H}}^2 = \theta_u^2 s/(1-s)/2$. For $\chi \to \infty$ we return to a uniform scattering material distribution. It is important that if the scattering screen is located at the distance $s = 2/3$, it produces the same size of scattering disk as the uniformly distributed medium, and in the absence of additional data we cannot distinguish these two cases.

We can solve Equation (3) for s

$$s = \frac{1 \pm \left[1 - \frac{(1+2r)(1-r)}{3r^2}\chi\right]^{1/2}}{2 + r^{-1}}, \qquad (4)$$

where $r = \theta_{\mathrm{H}}^2/\theta_u^2$ is the relative size of the scattering disk. Figure 1 shows the fractional distance s plotted against the relative power of scattering χ. For $r \in [0;1)$ there are solutions with $\chi \in [0; \chi_{\max}]$. Here $\chi_{\max} = 3r^2\left[(1+2r)(1-r)\right]^{-1}$ is the greatest possible contribution of the uniform medium along the line of sight in the total scattering for measured values of θ_{H} and τ_{sc}. If the influence of the uniform medium is negligible ($\chi = 0$), the position of the scattering screen is the farthest from the pulsar and corresponds to the single screen case: $s_0 = (1 + (2r)^{-1})^{-1}$. But if the uniform medium also contributes to the scattering, the screen stands closer to the pulsar. When the extended medium scatters the radio emission with maximum strength ($\chi = \chi_{\max}$), the distance to the thin screen is decreased down to $s_m = s_0/2$. Significantly, Equation (4) provides two possible values for s, even when the extended medium is absent. One of them is s_0 and the other one is 0. If χ is not 0, the first value of s is less than s_0 and it decreases with increasing χ. At the same time, the second possible value of s is increasing with χ up to s_m.

So if we take into account the uniform medium, we can place the scattering screen anywhere between its location corresponding to the case without a uniform medium and the pulsar. Moreover, we cannot clearly distinguish two possible screen positions even if the value of χ is known.

For $r > 1$ there is the only possible value of s that equals s_0 in the single screen model, which increases to
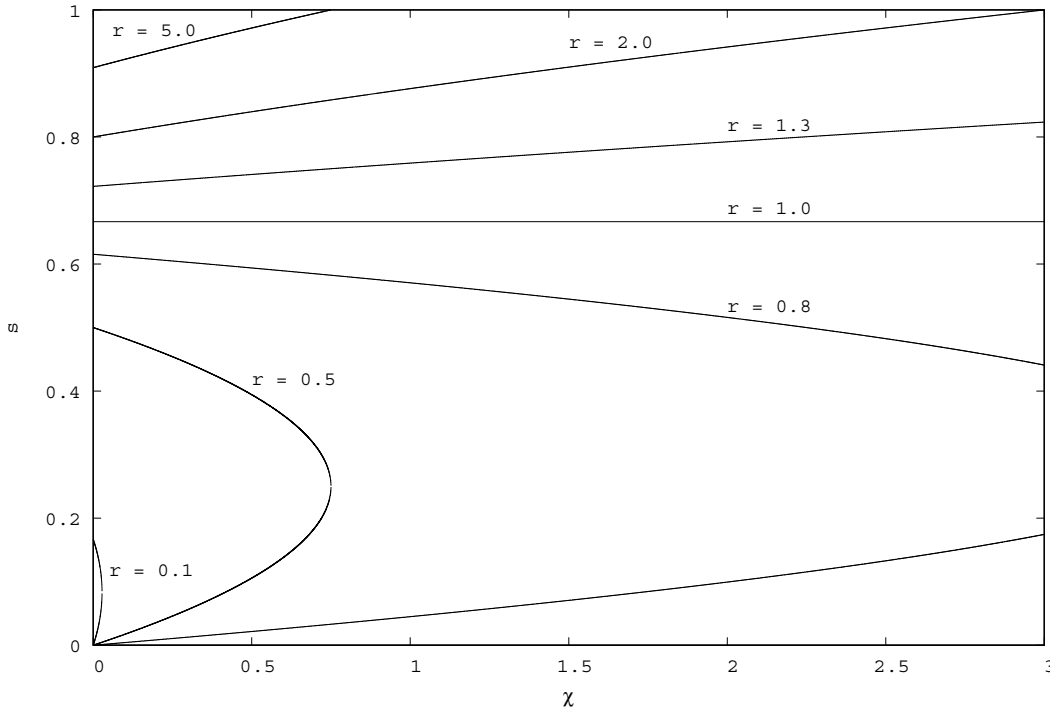
Fig. 1 The dependence of fractional distance s of the scattering screen plotted with the contribution of uniform medium χ for different measured relative sizes of scattering disk r. There are two groups of solutions. If the screen position in the uniform medium free model $s(\chi = 0) = s_0$ is less than $2/3$, the real position can be at any distance from pulsar to s_0. The maximum contribution of the uniform scattering medium occurs for the distance $s_0/2$. For screens with s_0 more than $2/3$, the real screen position can be anywhere between s_0 and the observer.

1 with the scattering strength of the uniform medium (χ). Since for $r > 1$ and $s > 2/3$, if we find out the screen distance s_0 is more than $2/3$, the presence of additional uniformly distributed material shifts the position of the scattering screen closer to the observer.

Thus we can locate the position of the scattering screen with certainty when it is close to the pulsar or the observer. In other cases, the true position of the screen can noticeably differ from the derived one for the single screen model.

2.3 Two Scattering Screens

Putney & Stinebring (2006) have shown that the secondary spectra of some pulsars exhibit multiple scintillation arcs with different curvatures. In our paper (Popov et al. 2016), we reported the detection of two parabolic arcs in the spectrum of pulsar B1933+16. Each single arc is produced by a separate thin screen whose location can be determined if the pulsar distance and its proper motion are known. For two scattering screens and uniform medium,

$$\psi(z) = \Psi_0 + \Psi_1 \delta(z - (D - d_1)) \\ + \Psi_2 \delta(z - (D - d_2)),$$

and ratios $\chi_1 = \Psi_0 D / \Psi_1$ and $\chi_2 = \Psi_0 D / \Psi_2$ of scattering strength for uniform medium and screens are bounded by the following relation

$$\chi_1 = \frac{3s_1 \left[2r(1 - s_1) - s_1 \right]}{1 - r - 3s_2 \left[2r(1 - s_2) - s_2 \right] / \chi_2}. \quad (5)$$

If determination of screen positions by means of scintillation arcs is accompanied by measurements of angular and temporal broadening of pulsars, we can estimate the relative scattering power of thin screens and the uniform medium. Neglecting the uniform medium, we obtain ratio $\chi_{12} = \Psi_1 / \Psi_2$ as a function of r, s_1, and s_2

$$\chi_{12} = \frac{s_2}{s_1} \frac{2(1 - s_2)r - s_2}{s_1 - 2(1 - s_1)r}. \quad (6)$$

It must be emphasized that observable angular and temporal broadening is produced by all screens and the uniform medium. The measurement of only θ_H and τ_{sc} does not make it possible to choose the single-component or the multi-component interstellar medium model.

3 COMPARISON WITH OBSERVATIONS

We estimated the position of thin scattering screens in the directions of three pulsars: B1641–45, B1749–28 and B1933+16 (Popov et al. 2016). For B1641–45, our estimates of s were 0.36 ± 0.02 for a distance to the pulsar of 4.5 ± 0.4 kpc. According to Equation (4), we can infer that the scattering screen is located somewhere between $s = 0$ and $s = 0.36$, i.e. the screen distance varies from 2.7 to 4.9 kpc depending on the pulsar distance. The model allows the maximum value of $\chi_{max} = 0.20 \pm 0.03$, i.e. a thin screen scatters radio waves at least five times more strongly than a uniform medium. The H II region G339.1–0.4 lies close to the line of sight at distance 3.3 kpc. If we assume that this H II region acts as the scattering screen, we can estimate the contribution of uniform medium χ as 0.14.

The distance to the pulsar B1749–28 is not well-known: $0.2^{+1.1}_{-0.1}$ kpc (Verbiest et al. 2012). At the lower limit of this distance, the contribution of the uniform medium is negligible, and the scattering screen is located very close to the pulsar. For the upper limit, the maximum contribution of a uniform medium still remains small but the screen can be shifted up to 300 pc from the pulsar.

In the secondary spectrum of B1933+16, there is a complex pattern that includes at least two parabolic arcs. Each arc corresponds to one scattering screen. We have located these screens at distances $1.0 - 1.1$ kpc and $2.4 - 3.7$ kpc from the observer for pulsar distance $3.7^{+1.3}_{-0.8}$ (Verbiest et al. 2012). The analysis of temporal and angular broadening gives the screen position $2.3 - 3.4$ kpc in the single screen approach. Neglecting the uniform medium described by Equation (6), the more distant screen scatters radio waves from 20 to 40 times more strongly than the closer one. This result was expected because the position of the more distant and powerful screen is in good agreement with the one scattering screen model. Equation (5) allows the existence of a uniform medium that scatters 10–30 times ($\chi_2 = 0.03 - 0.09$) less strongly than the more powerful screen. At higher values of χ_2, the weaker screen cannot exist.

4 CONCLUSIONS

We have shown that uniformly distributed scattering material affects the location of the thin scattering screen.

Without additional data, the real position of the scattering screen remains unknown and lies between the position calculated without the uniform medium and the pulsar ($d_s > D/3$) or the observer ($d_s < D/3$).

The uniform medium is responsible for no more than 20% of all scattering in the directions of pulsars B1641–45, B1749–28 and B1933+16. In the direction of the pulsar B1933+16, material that causes scattering is concentrated in two thin screens, with one of them being at least 10 times more powerful than the other.

References

Andrianov, A. S., Smirnova, T. V., Shishov, V. I., Gwinn, C., & Popov, M. V. 2017, Astronomy Reports, 61, 513

Blandford, R., & Narayan, R. 1985, MNRAS, 213, 591

Britton, M. C., Gwinn, C. R., & Ojeda, M. J. 1998, ApJ, 501, L101

Gwinn, C. R., Bartel, N., & Cordes, J. M. 1993, ApJ, 410, 673

Gwinn, C. R., Popov, M. V., Bartel, N., et al. 2016, ApJ, 822, 96

Lazio, T. J. W., Ojha, R., Fey, A. L., et al. 2008, ApJ, 672, 115

Popov, M. V., Andrianov, A. S., Bartel, N., et al. 2016, Astronomy Reports, 60, 792

Popov, M. V., Bartel, N., Gwinn, C. R., et al. 2017, MNRAS, 465, 978

Putney, M. L., & Stinebring, D. R. 2006, Chinese Journal of Astronomy and Astrophysics Supplement, 6, 233

Rao, A. P., & Ananthakrishnan, S. 1984, Nature, 312, 707

Readhead, A. C. S., & Hewish, A. 1972, Nature, 236, 440

Samus, N. N., & Li, Y., RAA (Research in Astronomy and Astrophysics), 2018, 18, 88

Shishov, V. I., Smirnova, T. V., Gwinn, C. R., et al. 2017, MNRAS, 468, 3709

Stinebring, D. R., McLaughlin, M. A., Cordes, J. M., et al. 2001, ApJ, 549, L97

Verbiest, J. P. W., Weisberg, J. M., Chael, A. A., Lee, K. J., & Lorimer, D. R. 2012, ApJ, 755, 39